

The Age-Redshift Relationship of Old Passive Galaxies

Jun-Jie Wei^{1,2}, Xue-Feng Wu^{1,3,4}, Fulvio Melia^{1,5}, Fa-Yin Wang^{6,7}, and Hai Yu^{6,7}

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¹Purple Mountain Observatory, Chinese Academy of Sciences, Nanjing 210008, China; jjwei@pmo.ac.cn, xfwu@pmo.ac.cn.

²University of Chinese Academy of Sciences, Beijing 100049, China.

³Chinese Center for Antarctic Astronomy, Nanjing 210008, China.

⁴Joint Center for Particle, Nuclear Physics and Cosmology, Nanjing University-Purple Mountain Observatory, Nanjing 210008, China.

⁵Department of Physics, The Applied Math Program, and Department of Astronomy, The University of Arizona, AZ 85721, USA; fmelia@email.arizona.edu.

⁶School of Astronomy and Space Science, Nanjing University, Nanjing 210093, China.

⁷Key Laboratory of Modern Astronomy and Astrophysics (Nanjing University), Ministry of Education, Nanjing 210093, China.

ABSTRACT

We use 32 age measurements of passively evolving galaxies as a function of redshift to test and compare the standard model (Λ CDM) with the $R_h = ct$ Universe. We show that the latter fits the data with a reduced $\chi^2_{\text{dof}} = 0.435$ for a Hubble constant $H_0 = 67.2^{+4.5}_{-4.0} \text{ km s}^{-1} \text{ Mpc}^{-1}$. By comparison, the optimal flat Λ CDM model, with two free parameters (including $\Omega_m = 0.12^{+0.54}_{-0.11}$ and $H_0 = 94.3^{+32.7}_{-35.8} \text{ km s}^{-1} \text{ Mpc}^{-1}$), fits the age- z data with a reduced $\chi^2_{\text{dof}} = 0.428$. Based solely on their χ^2_{dof} values, both models appear to account for the data very well, though the optimized Λ CDM parameters are only marginally consistent with those of the concordance model ($\Omega_m = 0.27$ and $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$). Fitting the age- z data with the latter results in a reduced $\chi^2_{\text{dof}} = 0.523$. However, because of the different number of free parameters in these models, selection tools, such as the Akaike, Kullback and Bayes Information Criteria, favour $R_h = ct$ over Λ CDM with a likelihood of $\sim 66.5\% - 80.5\%$ versus $\sim 19.5\% - 33.5\%$. These results are suggestive, though not yet compelling, given the current limited galaxy age- z sample. We carry out Monte Carlo simulations based on these current age measurements to estimate how large the sample would have to be in order to rule out either model at a $\sim 99.7\%$ confidence level. We find that if the real cosmology is Λ CDM, a sample of ~ 45 galaxy ages would be sufficient to rule out $R_h = ct$ at this level of accuracy, while ~ 350 galaxy ages would be required to rule out Λ CDM if the real Universe were instead $R_h = ct$. This difference in required sample size reflects the greater number of free parameters available to fit the data with Λ CDM.

Subject headings: cosmology: theory, observations – early universe – galaxy: general

1. Introduction

In recent years, the cosmic evolution has been studied using a diversity of observational data, including cosmic chronometers (Melia & Maier 2013), Gamma-ray bursts (GRBs; Wei et al. 2013), high- z quasars (Melia 2013, 2014a), strong gravitational lenses (Wei et al. 2014; Melia et al. 2015), and Type Ia SNe (Wei et al. 2015). In particular, the predictions of Λ CDM have been compared with those of a cosmology we refer to as the $R_h = ct$ Universe (Melia 2007; Melia & Shevchuk 2012). In all such one-on-one comparisons completed thus far, model selection tools show that the data favour $R_h = ct$ over Λ CDM.

The $R_h = ct$ Universe is a Friedmann-Robertson-Walker (FRW) cosmology that has much in common with Λ CDM, but includes an additional ingredient motivated by several theoretical and observational arguments (Melia 2007; Melia & Abdelqader 2009; Melia & Shevchuk 2012; see also Melia 2012a for a more pedagogical treatment). Like Λ CDM, it adopts an equation of state $p = w\rho$, with $p = p_m + p_r + p_{de}$ (for matter, radiation, and dark energy, respectively) and $\rho = \rho_m + \rho_r + \rho_{de}$, but goes one step further by specifying that $w = (\rho_r/3 + w_{de}\rho_{de})/\rho = -1/3$ at all times. Here, p is the total pressure and ρ is the total energy density. One might come away with the impression that this equation of state cannot be consistent with that (i.e., $w = [\rho_r/3 - \rho_\Lambda]/\rho$) in the standard model. But in fact if we ignore the constraint $w = -1/3$ and instead proceed to optimize the parameters in Λ CDM by fitting the data, the resultant value of w averaged over a Hubble time is actually $-1/3$ within the measurement errors (Melia 2007; Melia & Shevchuk 2012). In other words, though $w = (\rho_r/3 - \rho_\Lambda)/\rho$ in Λ CDM may be different from $-1/3$ from one moment to the next, its value averaged over the age of the Universe equals what it would have been in $R_h = ct$ all along (Melia 2015).

In this paper, we continue to compare the predictions of Λ CDM with those in the $R_h = ct$ Universe, this time focusing on the age-redshift relationship, which differs from

one expansion scenario to another. Though the current age of the universe may be similar in these two cosmologies, the age versus redshift relationship is not, particularly at high redshifts (Melia 2013, 2014b). Previous work with the age estimates of distant objects has already provided effective constraints on cosmological parameters (see, e.g., Alcaniz & Lima 1999; Lima & Alcaniz 2000; Jimenez & Loeb 2002; Jimenez et al. 2003; Capozziello et al. 2004; Friaça et al. 2005; Simon et al. 2005; Jain & Dev 2006; Pires et al. 2006; Dantas et al. 2007, 2009, 2011; Samushia et al. 2010). For example, using the simple criterion that the age of the Universe at any given redshift should always be greater than or equal to the age of the oldest object(s) at that redshift, the measured ages were used to constrain parameters in the standard model (Alcaniz & Lima 1999; Lima & Alcaniz 2000; Jain & Dev 2006; Pires et al. 2006; Dantas et al. 2007, 2011). Cosmological parameters have also been constrained by the measurement of the differential age $\Delta z/\Delta t$, where Δz is the redshift separation between two passively evolving galaxies having an age difference Δt (Jimenez & Loeb 2002; Jimenez et al. 2003). And the lookback time versus redshift measurements for galaxy clusters and passively evolving galaxies have been used to constrain dark energy models (Capozziello et al. 2004; Simon et al. 2005; Dantas et al. 2009; Samushia et al. 2010). This kind of analysis is therefore particularly interesting and complementary to those mentioned earlier, which are essentially based on distance measurements to a particular class of objects or physical rulers (see Jimenez & Loeb 2002, for a discussion on cosmological tests based on relative galaxy ages).

Age measurements of high- z objects have been valuable in constraining the cosmological parameters of the standard model even before dark energy was recognized as an essential component of the cosmic fluid (see, e.g., Bolte & Hogan 1995; Krauss & Turner 1995; Dunlop et al. 1996; Alcaniz & Lima 1999; Jimenez & Loeb 2002). Of direct relevance to the principal aim of this paper is the fact that, although the distance–redshift relationship is very similar in Λ CDM and the $R_h = ct$ Universe (even out to $z \gtrsim 6 - 7$; Melia 2012b; Wei et

al. 2013), the age–redshift dependence is not. This difference is especially noticeable in how we interpret the formation of structure in the early Universe. For example, the emergence of quasars at $z \gtrsim 6$, which are now known to be accreting at, or near, their Eddington limit (see, e.g., Willott et al. 2010; De Rosa et al. 2011). This presents a problem for Λ CDM because it is difficult to understand how $\sim 10^9 M_\odot$ supermassive black holes could have appeared only 700–900 Myr after the big bang. Instead, in $R_h = ct$, their emergence at redshift ~ 6 corresponds to a cosmic age of $\gtrsim 1.6$ Gyr, which was enough time for them to begin growing from $\sim 5 - 20 M_\odot$ seeds (presumably the remnants of Pop II and III supernovae) at $z \lesssim 15$ (i.e., *after* the onset of re-ionization) and still reach a billion solar masses by $z \sim 6$ via standard, Eddington-limited accretion (Melia 2013).

In this paper, we will broaden the base of support for this cosmic probe by demonstrating its usefulness in testing *competing* cosmological models. Following the methodology presented in Dantas et al. (2011), we will use 32 age measurements of passively evolving galaxies as a function of redshift (in the range $0.117 \leq z \leq 1.845$) to test the predicted age–redshift relationship of each model. From an observational viewpoint, because the age of a galaxy must be younger than the age of the Universe at any given redshift, there must be an incubation time, or delay factor τ , for the galaxy to form after the big bang. In principle, there could be a different τ_i for each object i since galaxies can form at different epochs. However, the simplest approach we can take is to begin with the assumption made in earlier work (see, e.g., Dantas et al. 2009, 2011; Samushia et al. 2010), i.e., we will adopt an average delay factor $\langle \tau \rangle$ and use it uniformly for every galaxy. But we shall also consider cases in which the delay factors τ_i are distributed, and study the impact of this non-uniformity on the overall fits to the data.

We will demonstrate that the current sample of galaxy ages favours the $R_h = ct$ Universe with a likelihood of $\sim 66.5 - 80.5\%$ of being correct, versus $\sim 19.5 - 33.5\%$

for Λ CDM. Though this result is still only marginal, it nonetheless calls for a significant increase in the sample of suitable galaxy ages in order to carry out more sophisticated and higher precision measurements. We will therefore also construct mock catalogs to investigate how big the sample of measured galaxy ages has to be in order to rule out one (or more) of these models.

The outline of this paper is as follows. In § 2, we will briefly describe the age-redshift test, and then constrain the cosmological parameters—both in the context of Λ CDM and the $R_h = ct$ universe, first using a uniform (average) delay factor $\langle\tau\rangle$ (§ 3), and then a distribution of τ_i values (§ 4). In § 5, we will discuss the model selection tools we use to test Λ CDM and the $R_h = ct$ cosmologies. In § 6, we will estimate the sample size required from future age measurements to reach likelihoods of $\sim 99.7\%$ and $\sim 0.3\%$ (i.e., 3σ confidence limits) when using model selection tools to compare these two models, and we will end with our conclusions in § 7.

2. The age-redshift test

The theoretical age of an object at redshift z is given as

$$t^{\text{th}}(z, \mathbf{p}) = \int_z^\infty \frac{dz'}{(1+z')H(z', \mathbf{p})} , \quad (1)$$

where \mathbf{p} stands for all the parameters of the cosmological model under consideration and $H(z, \mathbf{p})$ is the Hubble parameter at redshift z . From an observational viewpoint, the total age of a given object (e.g., a galaxy) at redshift z is given by $t^{\text{obs}}(z) = t_G(z) + \tau$, where $t_G(z)$ is the estimated age of its oldest stellar population and τ is the incubation time, or delay factor, which accounts for our ignorance about the amount of time elapsed since the big bang to the epoch of star creation.

To compute model predictions for the age $t^{\text{th}}(z, \mathbf{p})$ in Equation (1), we need an

expression for $H(z, \mathbf{p})$. As we have seen, Λ CDM assumes specific constituents in the density, written as $\rho = \rho_r + \rho_m + \rho_{de}$. These densities are often written in terms of today's critical density, $\rho_c \equiv 3c^2 H_0^2 / 8\pi G$, represented as $\Omega_m \equiv \rho_m / \rho_c$, $\Omega_r \equiv \rho_r / \rho_c$, and $\Omega_{de} \equiv \rho_{de} / \rho_c$. H_0 is the Hubble constant. In a flat universe with zero spatial curvature, the total scaled energy density is $\Omega \equiv \Omega_m + \Omega_r + \Omega_{de} = 1$. When dark energy is included with an unknown equation-of-state, $p_{de} = w_{de}\rho_{de}$, the most general expression for the Hubble parameter is

$$H(z, \mathbf{p}) = H_0 \left[\Omega_m(1+z)^3 + \Omega_r(1+z)^4 + \Omega_k(1+z)^2 + \Omega_{de}(1+z)^{3(1+w_{de})} \right]^{1/2}, \quad (2)$$

where Ω_k is defined similarly to Ω_m and represents the spatial curvature of the Universe. Of course, Ω_r ($\sim 5 \times 10^{-5}$) is known from the current temperature ($\simeq 2.725$ K) of the CMB, and the value of H_0 . In addition, we assume a flat Λ CDM cosmology with $\Omega_k = 0$, for which $\Omega_{de} = 1 - \Omega_m - \Omega_r$, thus avoiding the introduction of Ω_{de} as an additional free parameter. So for the basic Λ CDM model, \mathbf{p} includes three free parameters: Ω_m , w_{de} , and H_0 , though to be as favorable as possible to this model, we will also assume that dark energy is a cosmological constant, with $w_{de} = -1$, leaving only two adjustable parameters. However, when we consider the concordance model with the assumption of prior parameter values, the fits have zero degrees of freedom from the cosmology itself.

The $R_h = ct$ Universe is a flat Friedmann-Robertson-Walker (FRW) cosmology that strictly adheres to the constraints imposed by the simultaneous application of the cosmological principle and Weyl's postulate (Melia 2012b; Melia & Shevchuk 2012). When these ingredients are applied to the cosmological expansion, the gravitational horizon $R_h = c/H$ must always be equal to ct . This cosmology is therefore very simple, because $a(t) \propto t$, which also means that $1+z = 1/t$, with the (standard) normalization that $a(t_0) = 1$. Therefore, in the $R_h = ct$ Universe, we have the straightforward scaling

$$H(z, \mathbf{p}) = (1+z)H_0. \quad (3)$$

Notice, in particular, that the expansion rate $H(z)$ in this model has only one free

parameter, i.e., \mathbf{p} is H_0 . From Equation (3), the age of the $R_h = ct$ Universe at redshift z is simply

$$t^{\text{th}}(z, H_0) = \frac{1}{(1+z)H_0}. \quad (4)$$

To carry out the age-redshift analysis of Λ CDM and $R_h = ct$, we will first attempt to fit the ages of 32 old passive galaxies distributed over the redshift interval $0.117 \leq z \leq 1.845$ (Simon et al. 2005), listed in Table 1 of Samushia et al. (2010), assuming a uniform value of the time delay τ for every galaxy. Following these authors, we will also assume a 12% one-standard deviation uncertainty on the age measurements (Dantas et al. 2009, 2011; Samushia et al. 2010). The total sample is composed of three sub-samples: 10 field early-type galaxies from Treu et al. (1999, 2001, 2002), whose ages were obtained by using the SPEED models of Nolan et al. (2001) and Jimenez et al. (2004); 20 red galaxies from the publicly released Gemini Deep Deep Survey (GDDS), whose integrated light is fully dominated by evolved stars (Abraham et al. 2004; McCarthy et al. 2004); and the 2 radio galaxies LBDS 53W091 and LBDS 53W069 (Dunlop et al. 1996; Spinrad et al. 1997). These data were first collated by Jimenez et al. (2003).

3. Optimization of the Model Parameters Using a Uniform τ

In subsequent sections of this paper, we will study the impact of a distributed incubation time on the overall fits to the data. However, in this first simple approach (see, e.g., Dantas et al. 2009, 2011; Samushia et al. 2010) we will assume an average delay factor $\langle\tau\rangle$ and use it uniformly for every galaxy, so that we may compare our results to those of previous work.

For each model, we optimize the fit by finding the set of parameters (\mathbf{p}) that minimize

the χ^2 , using the statistic

$$\begin{aligned}\chi_{age}^2(\tau, \mathbf{p}) &= \sum_{i=1}^{32} \frac{[t^{\text{th}}(z_i, \mathbf{p}) - t_G(z_i) - \langle \tau \rangle]^2}{\sigma_{t_{G,i}}^2} \\ &\equiv A - 2 * \langle \tau \rangle * B + \langle \tau \rangle^2 * C ,\end{aligned}\tag{5}$$

where $A \equiv \sum [t^{\text{th}}(z_i, \mathbf{p}) - t_G(z_i)]^2 / \sigma_{t_{G,i}}^2$, $B \equiv \sum [t^{\text{th}}(z_i, \mathbf{p}) - t_G(z_i)] / \sigma_{t_{G,i}}^2$, and $C \equiv \sum 1 / \sigma_{t_{G,i}}^2$. The dispersions $\sigma_{t_{G,i}}$ represent the uncertainties on the age measurements of the sample galaxies. Given the form of Equation (5), we can marginalize $\langle \tau \rangle$ by minimizing χ_{age}^2 , which has a minimum at $\langle \tau \rangle = B/C$, with a value $\hat{\chi}_{age}^2 = A - B^2/C$. Note that this procedure allows us to determine the optimized value of $\langle \tau \rangle$ along with the best-fit parameters of the model being tested.

3.1. Λ CDM

In the concordance Λ CDM model, the dark-energy equation of state parameter, w_{de} , is exactly -1 . The Universe is flat, $\Omega_{\text{de}} = 1 - \Omega_{\text{m}} - \Omega_{\text{r}}$, so there remain only two free parameters: Ω_{m} and H_0 . Type Ia SN measurements (see, e.g., Garnavich et al. 1998; Perlmutter et al. 1998, 1999; Riess et al. 1998; Schmidt et al. 1998), CMB anisotropy data (see, e.g., Ratra et al. 1999; Podariu et al. 2001; Spergel et al. 2003; Komatsu et al. 2009, 2011; Hinshaw et al. 2013), and baryon acoustic oscillation (BAO) peak length scale estimates (see, e.g., Percival et al. 2007; Gaztañaga et al. 2009; Samushia & Ratra 2009), strongly suggest that we live in a spatially flat, dark energy-dominated universe with concordance parameter values $\Omega_{\text{m}} \approx 0.27$ and $H_0 \approx 70.0 \text{ km s}^{-1} \text{ Mpc}^{-1}$.

In order to gauge how well Λ CDM and $R_{\text{h}} = ct$ account for the galaxy age-redshift measurements, we will first attempt to fit the data with this concordance model, using prior values for all the parameters but one, i.e., the unknown average delay time $\langle \tau \rangle$. We will improve the fit as much as possible by marginalizing $\langle \tau \rangle$, as described above. Fitting

the 32 age-redshift measurements with a theoretical $t^{\text{th}}(z)$ function using $\Omega_{\text{m}} = 0.27$ and $H_0 = 70.0 \text{ km s}^{-1} \text{ Mpc}^{-1}$, we obtain an optimized delay factor $\langle\tau\rangle = 1.36 \text{ Gyr}$, and a $\chi^2_{\text{dof}} = 16.17/31 = 0.523$, remembering that all of the ΛCDM parameters are assumed to have prior values, except for $\langle\tau\rangle$.

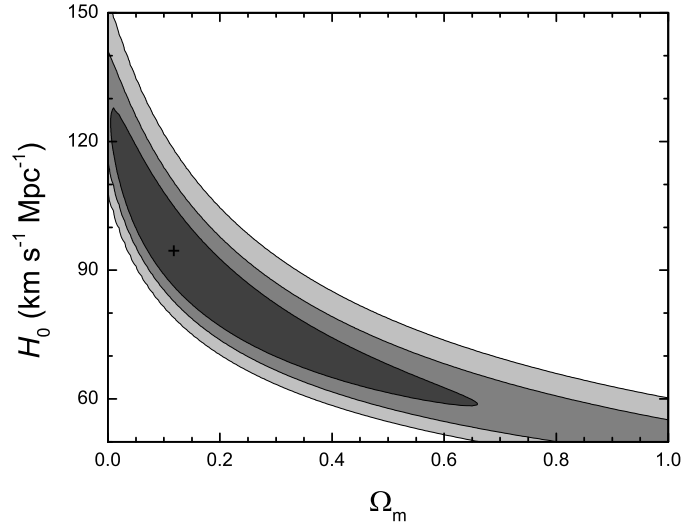


Fig. 1.— 1-3 σ -constraint contours for the flat ΛCDM model, using the 32 age-redshift data. The cross indicates the best-fit pair $(\Omega_{\text{m}}, H_0) = (0.12, 94.3)$.

If we relax the priors, and allow both Ω_{m} and H_0 to be free parameters, we obtain best-fit values $(\Omega_{\text{m}}, H_0) = (0.12, 94.3 \text{ km s}^{-1} \text{ Mpc}^{-1})$, as illustrated in Figure 1. The delay factor corresponding to this best fit is $\langle\tau\rangle = 1.62 \text{ Gyr}$, with a $\chi^2_{\text{dof}} = 12.42/29 = 0.428$. Figure 1 also shows the 1-3 σ constraint contours of the probability function in the $\Omega_{\text{m}}\text{-}H_0$ plane. Insofar as the ΛCDM model is concerned, this optimization has improved the quality of the fit as gauged by the reduced χ^2_{dof} , though the parameter values are quite different from those of the concordance model. Nonetheless, these two sets of values are still marginally consistent with each other because the data are not good enough yet to improve the precision with which Ω_{m} and H_0 are determined. It is also possible that treating $\langle\tau\rangle$ as

a uniform variable for all galaxies may be over-constraining, but we will relax this condition in subsequent sections and consider situations in which τ_i may be different for each galaxy in the sample. We shall see that these optimized values change quantitatively, though the qualitative results and conclusions remain the same. The contours in Figure 1 show that at the 1σ -level, we have $58.5 < H_0 < 127.0 \text{ km s}^{-1} \text{ Mpc}^{-1}$, and $0.01 < \Omega_m < 0.66$. The cross indicates the best-fit pair. For the sake of a direct one-on-one comparison between Λ CDM and the $R_h = ct$ Universe, the current status with these data therefore suggests that we should use the concordance parameter values, which are supported by many other kinds of measurements, as described above.

3.2. The $R_h = ct$ Universe

Regardless of what constituents may be present in the cosmic fluid, insofar as the expansion dynamics is concerned, the $R_h = ct$ Universe always has just one free parameter, H_0 . The results of fitting the age-redshift data with this cosmology are shown in Figure 2 (solid line). We see here that the best fit corresponds to $H_0 = 67.2^{+4.5}_{-4.0} \text{ km s}^{-1} \text{ Mpc}^{-1}$ (1σ). The corresponding delay factor is $\langle\tau\rangle = 2.72 \text{ Gyr}$. With $32 - 2 = 30$ degrees of freedom, we have $\chi^2_{\text{dof}} = 13.05/30 = 0.435$.

To facilitate a direct comparison between Λ CDM and $R_h = ct$, we show in Figure 3 the galaxy ages (i.e., $t_G + \langle\tau\rangle$), together with the best-fit theoretical curves for the $R_h = ct$ Universe (with $H_0 = 67.2 \text{ km s}^{-1} \text{ Mpc}^{-1}$ and $\langle\tau\rangle = 2.72 \text{ Gyr}$), the concordance model (with $\langle\tau\rangle = 1.36 \text{ Gyr}$ and prior values for all the other parameters), and for the optimized Λ CDM model (with $H_0 = 94.3 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_m = 0.12$, and $\langle\tau\rangle = 1.62 \text{ Gyr}$). As described above, $\langle\tau\rangle$ is the average incubation time or delay factor, which accounts for our ignorance concerning the amount of time elapsed since the big bang to the initial formation of the object. At the very minimum, $\langle\tau\rangle$ must be greater than $\sim 300 \text{ Myr}$, this being the

time at which Population III stars would have established the necessary conditions for the subsequent formation of Population II stars (see, e.g., Melia 2014b and references cited therein). The galaxies could not have formed any earlier than this, based on the physics we know today.

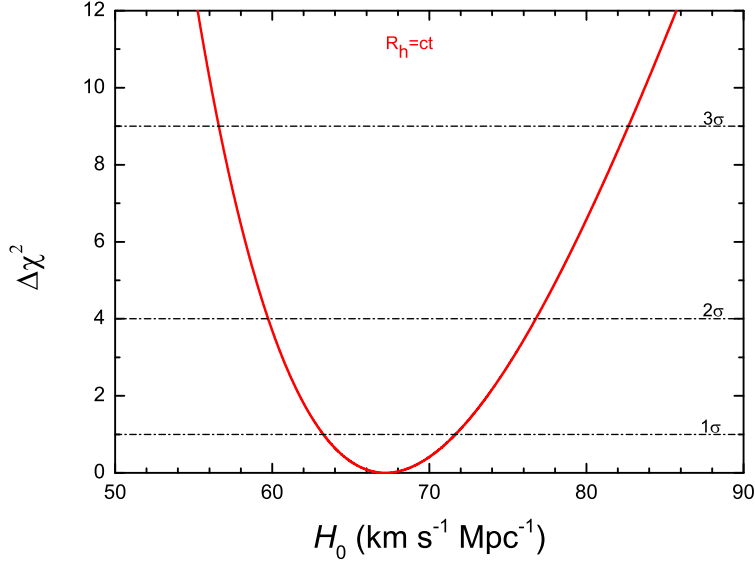


Fig. 2.— Constraints on the Hubble constant, H_0 , in the context of $R_h = ct$.

In this figure, we also show $t_G + 300$ Myr versus z (circles) to illustrate the minimum possible ages the galaxies could have, given what we now know about the formation of Population II and Population III stars. We note that the best fit value of $\langle\tau\rangle$, in both Λ CDM and $R_h = ct$, is fully consistent with the supposition that all of the galaxies should have formed after the transition from Population III to Population II stars at $t \sim 300$ Myr. Strictly based on their χ^2_{dof} values, the concordance Λ CDM model and the $R_h = ct$ Universe appear to fit the passive galaxy age-redshift relationship (i.e., $t_G + \langle\tau\rangle$ versus z) comparably well. However, because these models formulate their observables (such as the theoretical ages in Equations 1 and 4) differently, and because they do not have the same number of

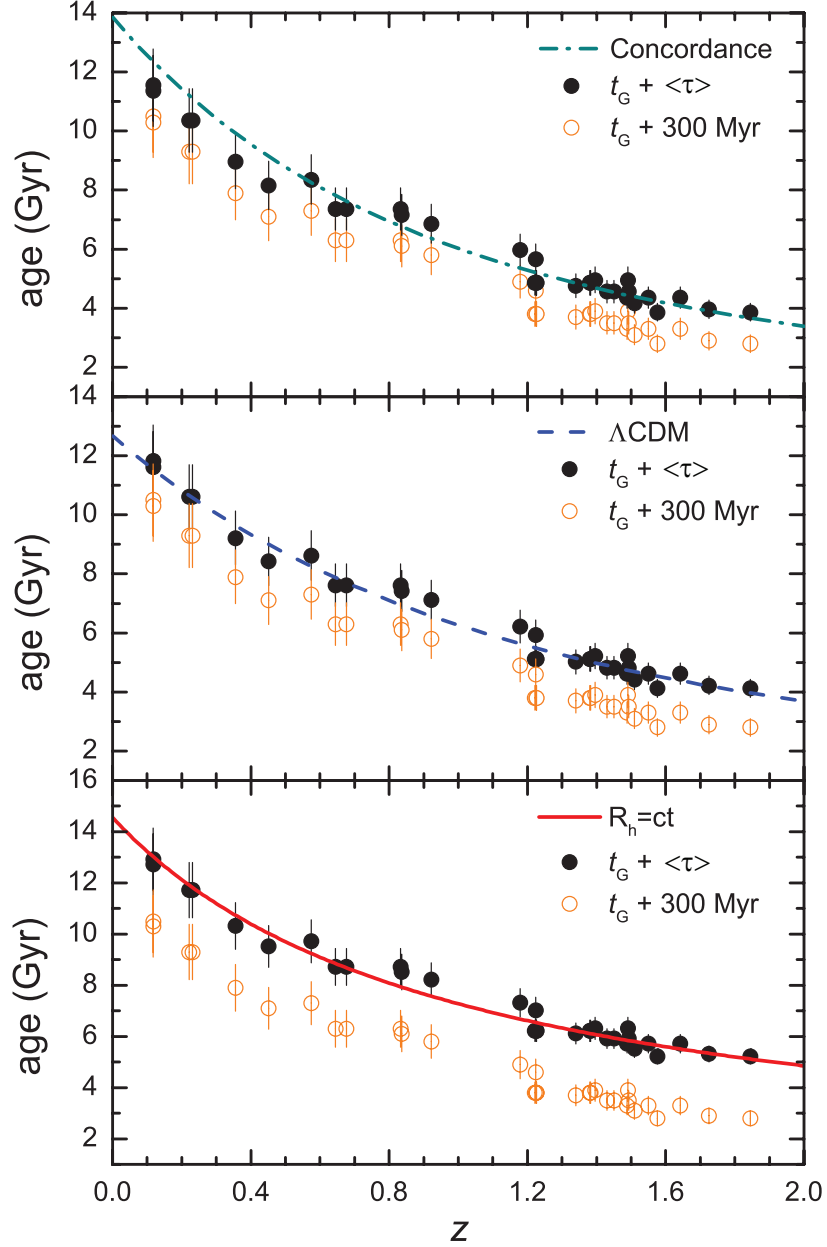


Fig. 3.— The complete age-redshift sample (*solid points*), and the best-fit theoretical curves: (*dot-dashed line*) the concordance model, with its sole optimized parameter $\langle \tau \rangle = 1.36 \text{ Gyr}$; (*dashed line*) the standard, flat Λ CDM cosmology, with optimized parameters $\Omega_m = 0.12$, $H_0 = 94.3^{+32.7}_{-35.8} (1\sigma) \text{ km s}^{-1} \text{ Mpc}^{-1}$, and $\langle \tau \rangle = 1.62 \text{ Gyr}$; (*solid line*) the $R_h = ct$ Universe, with $H_0 = 67.2^{+4.5}_{-4.0} (1\sigma) \text{ km s}^{-1} \text{ Mpc}^{-1}$ and $\langle \tau \rangle = 2.72 \text{ Gyr}$. The *empty circles* show the minimum ages the galaxies could have using $t_G(z_i) + 300 \text{ Myr}$.

free parameters, a comparison of the likelihoods for either being closer to the ‘true’ model must be based on model selection tools, which we discuss in § 5 below. But first, we will strengthen this analysis by considering possibly more realistic distributions of the delay time τ .

4. Optimization of the Model Parameters Using a Distributed τ

We now relax the constraint that the time delay should have the same value $\langle\tau\rangle$ for every galaxy, and instead consider two representative distributions: (i) a Gaussian

$$P(\tau) \propto \exp -\frac{(\tau - \tau_c)^2}{2\sigma_\tau^2} , \quad (6)$$

where the mean value τ_c is to be optimized for each theoretical fit, given some dispersion σ_τ ; and (ii) a top-hat

$$P(\tau) = \begin{cases} \text{const} & (\tau_c - \sigma_\tau < \tau < \tau_c + \sigma_\tau) \\ 0 & (\text{otherwise}) , \end{cases} \quad (7)$$

with σ_τ now representing the width of the distribution.

To isolate the various influences as much as possible, we begin with the concordance Λ CDM model, for which $\Omega_m = 0.27$ and $w_{de} = -1$ (i.e., dark energy is assumed to be a cosmological constant), though we optimize the Hubble constant to maximize the quality of the fit. For each distribution $P(\tau)$ and assumed value of σ_τ , we randomly assign the time delay τ_i to each galaxy and then find the best-fit values of H_0 and τ_c by minimizing the χ^2 using the statistic

$$\chi_{age}^2(\tau_c, \mathbf{p}) = \sum_{i=1}^{32} \frac{[t^{\text{th}}(z_i, \mathbf{p}) - t_G(z_i) - \tau_i(\tau_c, \sigma_\tau)]^2}{\sigma_{t_{G,i}}^2} . \quad (8)$$

The left-hand panels of Figure 4 show the ensuing distributions of τ_c and H_0 values for a Gaussian $P(\tau)$, and three different assumed dispersions σ_τ . In this case, the optimized

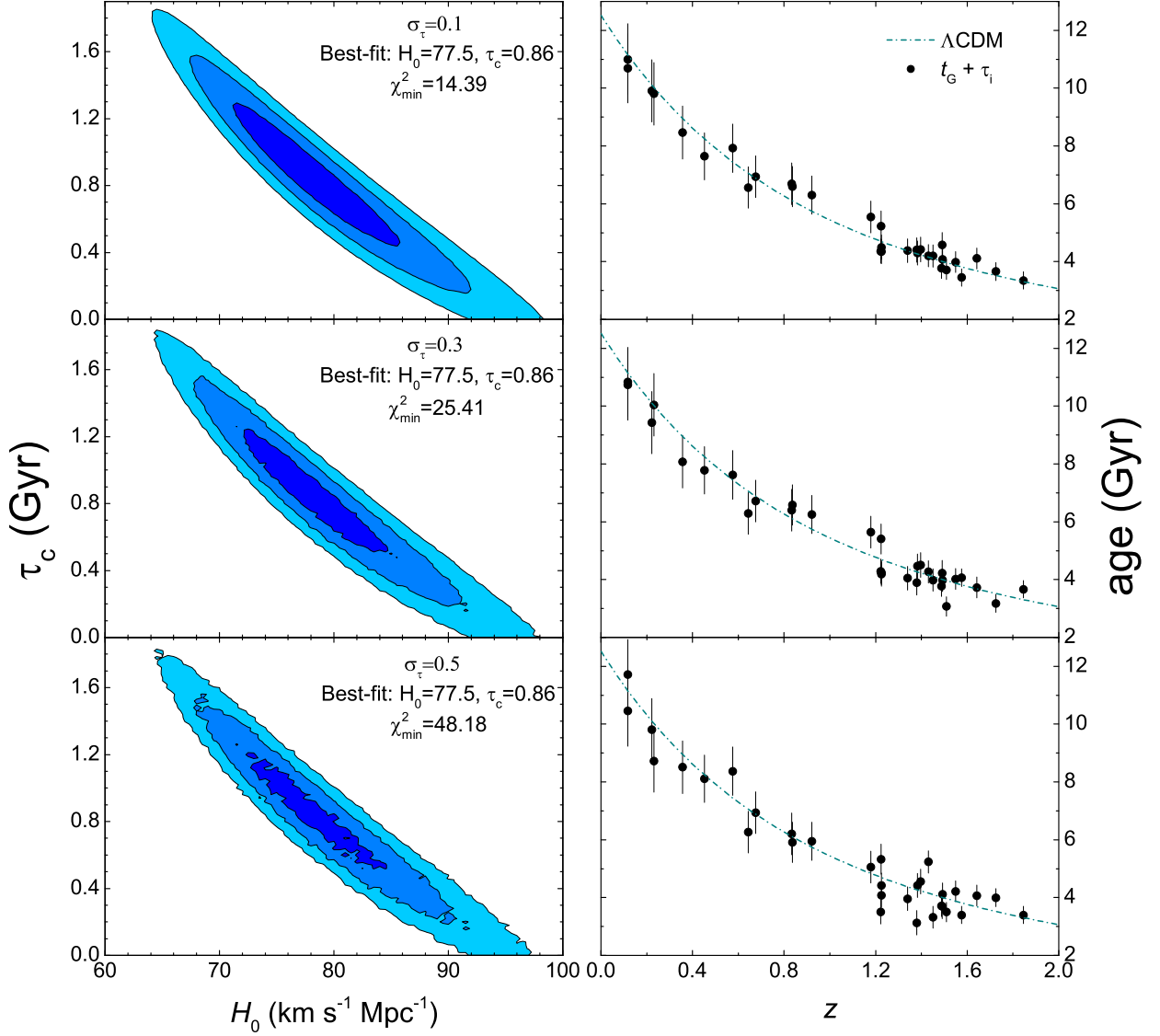


Fig. 4.— Λ CDM with a Gaussian distribution of τ values. Left-hand panels: fitted values of τ_c and H_0 using the complete age-redshift sample (right-hand panels; *solid points*). The theoretical curves (right-hand panel; *dot-dashed curves*) correspond to the parameter values that minimize the χ^2 (shown on the left). For Λ CDM, a Gaussian distribution in τ results in an optimized value of the Hubble constant (~ 77.5 km s⁻¹ Mpc⁻¹) only weakly dependent on σ_τ , and a mean delay time $\tau_c \sim 0.86$ Gyr.

Hubble constant ($\sim 77.5 \text{ km s}^{-1} \text{ Mpc}^{-1}$) is effectively independent of σ_τ , while the best-fit value of the mean delay time τ_c is restricted to $\sim 0.86 \text{ Gyr}$. Not surprisingly, the scatter about the best-fit theoretical curve worsens as σ_τ increases, resulting in larger values of χ_{\min}^2 .

Assuming a top-hat $P(\tau)$ with ΛCDM produces the results shown in Figure 5. The best-fit values of H_0 and τ_c are very similar to those associated with Figure 4. In this case, both H_0 and τ_c are effectively independent of the assumed distribution width σ_τ .

We follow the same procedure for the $R_h = ct$ Universe, first considering a Gaussian distribution of τ_i values (Figure 6), followed by the top-hat distribution (Figure 7). The comparison between these two cosmological models may be summarized as follows: the best-fit results are very similar for both the Gaussian and top-hat $P(\tau)$ distributions, for the same cosmological model; strictly based on their minimum χ^2 values, the concordance ΛCDM model and the $R_h = ct$ Universe appear to fit the passive galaxy age-redshift relationship (i.e., $t_G + \tau_i$ versus z) comparably well, independent of what kind of time-delay distribution is assumed.

5. Model Selection Tools

Several model selection tools used in cosmology (see, e.g., Melia & Maier 2013, and references cited therein) include the Akaike Information Criterion, $\text{AIC} = \chi^2 + 2n$, where n is the number of free parameters (Liddle 2007), the Kullback Information Criterion, $\text{KIC} = \chi^2 + 3n$ (Cavanaugh 2004), and the Bayes Information Criterion, $\text{BIC} = \chi^2 + (\ln N)n$, where N is the number of data points (Schwarz 1978). In the case of AIC, with AIC_α characterizing model \mathcal{M}_α , the unnormalized confidence that this model is true is the Akaike

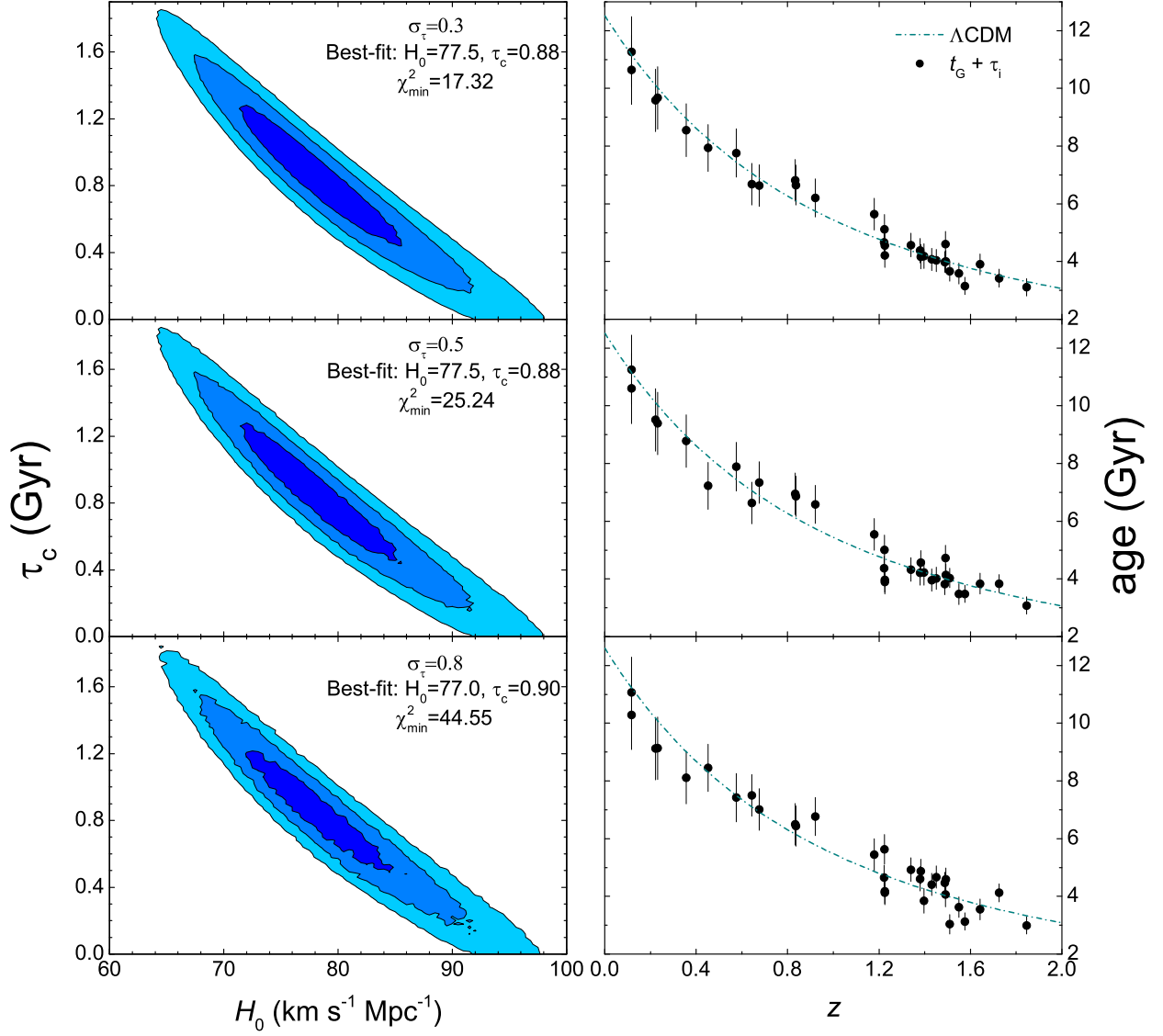


Fig. 5.— Same as Figure 4, except now for the top-hat distribution $P(\tau)$ given in Equation (7). In this case, both H_0 and τ_c are effectively independent of the assumed distribution width σ_τ .

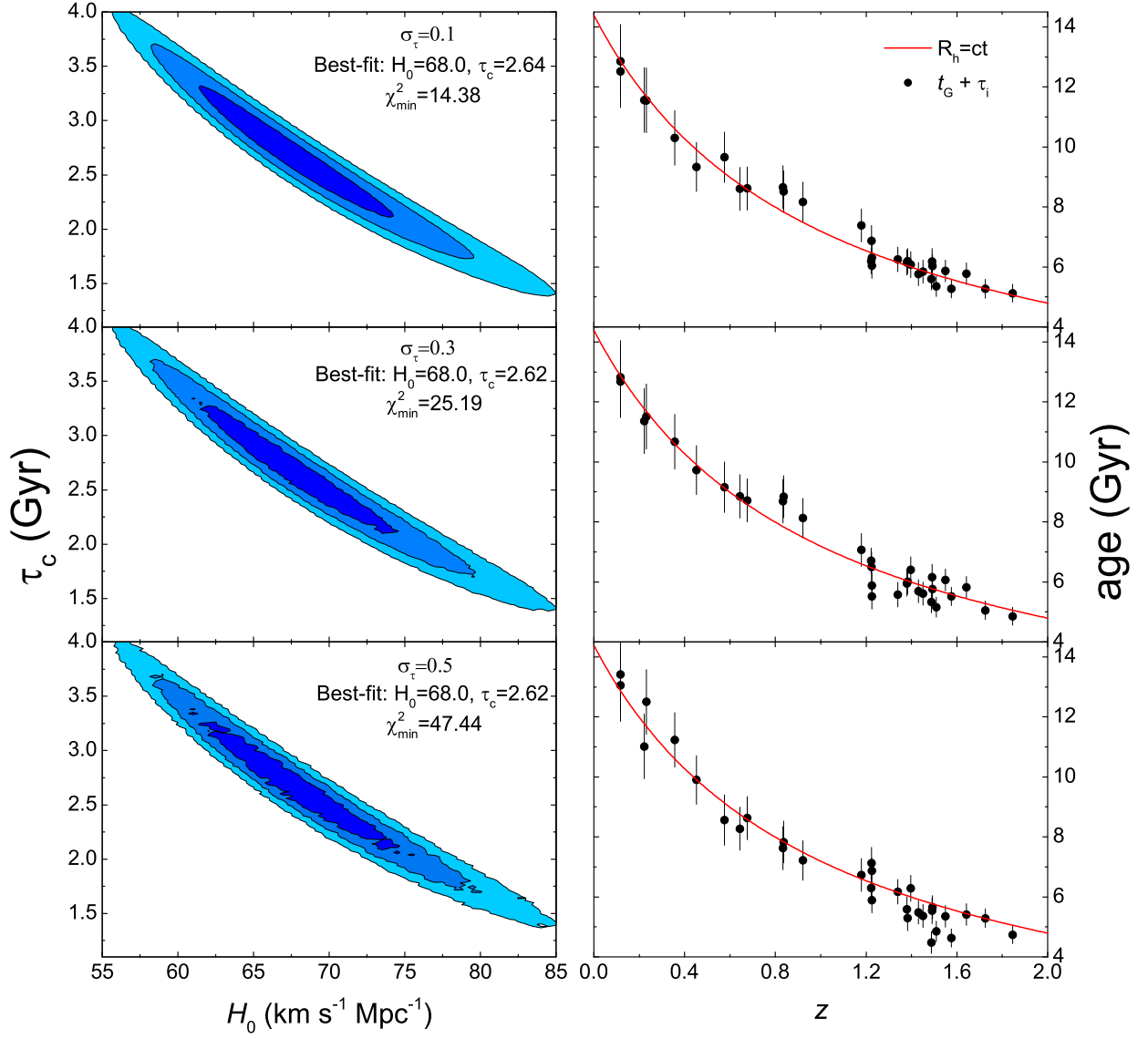


Fig. 6.— Same as Figure 4, except now for the $R_h = ct$ Universe.

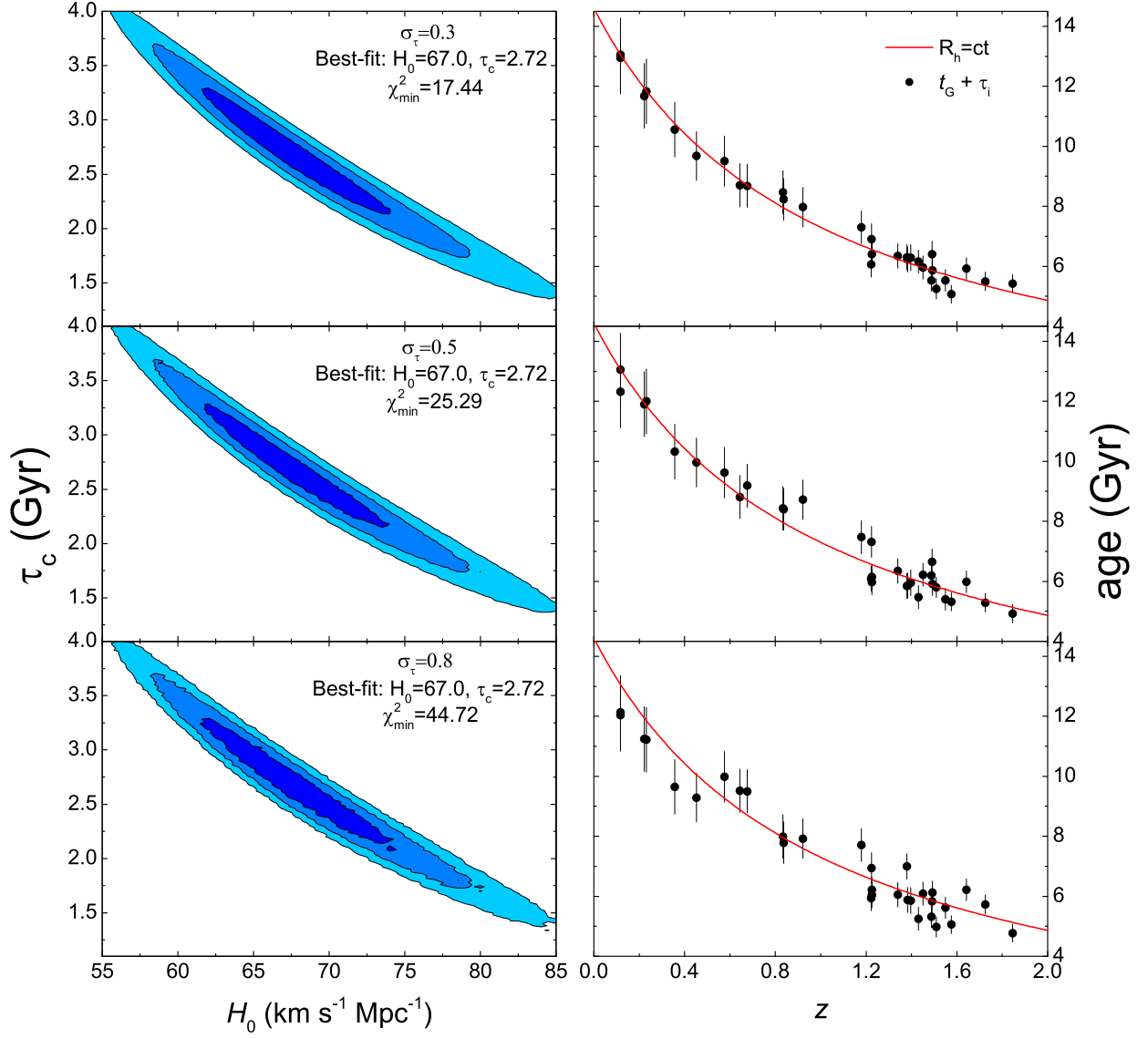


Fig. 7.— Same as Figure 6, except now for the top-hat distribution $P(\tau)$ given in Equation (7).

weight $\exp(-\text{AIC}_\alpha/2)$. Model \mathcal{M}_α has likelihood

$$P(\mathcal{M}_\alpha) = \frac{\exp(-\text{AIC}_\alpha/2)}{\exp(-\text{AIC}_1/2) + \exp(-\text{AIC}_2/2)} \quad (9)$$

of being the correct choice in this one-on-one comparison. Thus, the difference $\Delta\text{AIC} \equiv \text{AIC}_2 - \text{AIC}_1$ determines the extent to which \mathcal{M}_1 is favoured over \mathcal{M}_2 . For Kullback and Bayes, the likelihoods are defined analogously.

For the case of the average delay factor $\langle\tau\rangle$, with the optimized fits we have reported in this paper, our analysis of the age- z shows that the KIC does not favour either $R_h = ct$ or the concordance model when we assume prior values for all of its parameters. The calculated KIC likelihoods in this case are $\approx 51.5\%$ for $R_h = ct$, versus $\approx 48.5\%$ for ΛCDM . However, if we relax some of the priors, and allow both Ω_m and H_0 to be optimized in ΛCDM , then $R_h = ct$ is favoured over the standard model with a likelihood of $\approx 66.5\%$ versus 33.5% using AIC, $\approx 76.6\%$ versus $\approx 23.4\%$ using KIC, and $\approx 80.5\%$ versus $\approx 19.5\%$ using BIC.

For the distributed time delays discussed in Section 4, the model selection criteria result in the likelihoods shown in Table 1. Note that in this case, both the concordance ΛCDM and $R_h = ct$ models have the same free parameters (i.e., H_0 and τ_c), so the information criteria should all provide the same results. For the sake of clarity, we therefore show only the AIC results in Table 1, where we see that the AIC does not favour either $R_h = ct$ or the concordance model, regardless of which distribution is adopted for the incubation time.

6. Numerical Simulations

The results of our analysis suggest that the measurement of galaxy ages may be used to identify the preferred model in a one-on-one comparison. In using the model selection tools, the outcome $\Delta \equiv \text{AIC}_1 - \text{AIC}_2$ (and analogously for KIC and BIC) is judged ‘positive’

in the range $\Delta = 2 - 6$, and ‘strong’ for $\Delta > 6$. As we have seen, the adoption of prior values for the parameters in Λ CDM produces comparable likelihood outcomes for both models, regardless of whether we assume a uniform time delay $\langle\tau\rangle$, or a distribution of values. Of course, a proper statistical comparison between Λ CDM and $R_h = ct$ should not have to rely on prior values, particularly since the optimized cosmological parameters differ from survey to survey. If we don’t assume prior values for the parameters in Λ CDM, and optimize them to produce a best fit to the galaxy-age data, the corresponding Δ using the currently known 32 galaxy ages falls within the ‘positive’ range in favor of $R_h = ct$, though not yet the strong one. These results are therefore suggestive, but still not sufficient to rule out either model. In this section, we will therefore estimate the sample size required to significantly strengthen the evidence in favour of $R_h = ct$ or Λ CDM, by conservatively seeking an outcome even beyond $\Delta \simeq 11.62$, i.e., we will see what is required to produce a likelihood $\sim 99.7\%$ versus $\sim 0.3\%$, corresponding to 3σ .

Since the results do not appear to depend strongly on whether one chooses a uniform $\langle\tau\rangle$ for all the galaxies, or a distribution of individual τ_i values, we will first use the same average value $\langle\tau\rangle$ in our simulations, and then discuss how these results would change for a distributed incubation time. We will consider two cases: one in which the background cosmology is assumed to be Λ CDM, and a second in which it is $R_h = ct$, and we will attempt to estimate the number of galaxy ages required in each case in order to rule out the alternative (incorrect) model at a $\sim 99.7\%$ confidence level. The synthetic galaxy ages are each characterized by a set of parameters denoted as $(z, t[z])$, where $t(z) = t_G + \langle\tau\rangle$. We generate the synthetic sample using the following procedure:

1. Since the current 32 old passively evolving galaxies are distributed over the redshift interval $0.117 \leq z \leq 1.845$, we assign z uniformly between 0.1 and 2.0.
2. With the mock z , we first infer $t(z)$ from Equations (1) and (4) corresponding either

to a flat Λ CDM cosmology with $\Omega_{\text{m}} = 0.27$ and $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$ (§ 4.2), or the $R_{\text{h}} = ct$ Universe with $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$ (§ 4.1). We then assign a deviation (Δt) to the $t(z)$ value, i.e., we infer $t'(z)$ from a normal distribution whose center value is $t(z)$ and $\sigma = 0.35$ is its deviation (see Bengaly et al. 2014). The typical value of $\sigma = 0.35$ is taken from the current (observed) sample, which yields a mean and median deviation of $\sigma = 0.38$ and 0.33, respectively.

3. Since the observed error σ_t is about 12% of the age measurement, we will also assign a dispersion $\sigma_t = 0.12 t'(z)$ to the synthetic sample.

This sequence of steps is repeated for each galaxy in the sample, which is enlarged until the likelihood criterion discussed above is reached. As with the real 32-ages sample, we optimize the model fits by minimizing the χ^2 function $\chi^2 = \sum \left\{ [t^{\text{th}}(z_i, \mathbf{p}) - t'(z_i)]^2 / \sigma_{t,i}^2 \right\}$. This minimization is equivalent to maximizing the likelihood function $\mathcal{L} \propto \exp(-\chi^2/2)$. We employ Markov-chain Monte Carlo techniques. In each Markov chain, we generate 10^5 samples according to the likelihood function. Then we derive the cosmological parameters from a statistical analysis of the sample.

6.1. Assuming $R_{\text{h}} = ct$ as the Background Cosmology

We have found that a sample of at least 350 galaxy ages is required in order to rule out Λ CDM at the $\sim 99.7\%$ confidence level. The optimized parameters corresponding to the best-fit Λ CDM model for these simulated data are displayed in Figure 8. To allow for the greatest flexibility in this fit, we relax the assumption of flatness, and allow Ω_{de} to be a free parameter, along with Ω_{m} . Figure 8 shows the 1-D probability distribution for each parameter (Ω_{m} , Ω_{de} , H_0), and 2-D plots of the 1σ and 2σ confidence regions for two-parameter combinations. The best-fit values for Λ CDM using the simulated

Table 1. AIC Likelihood Estimation

Distributed τ	σ_τ	Λ CDM	$R_h = ct$
Gaussian	0.1	50%	50%
	0.3	47%	53%
	0.5	41%	59%
Top-hat	0.3	51%	49%
	0.5	50%	50%
	0.8	52%	48%

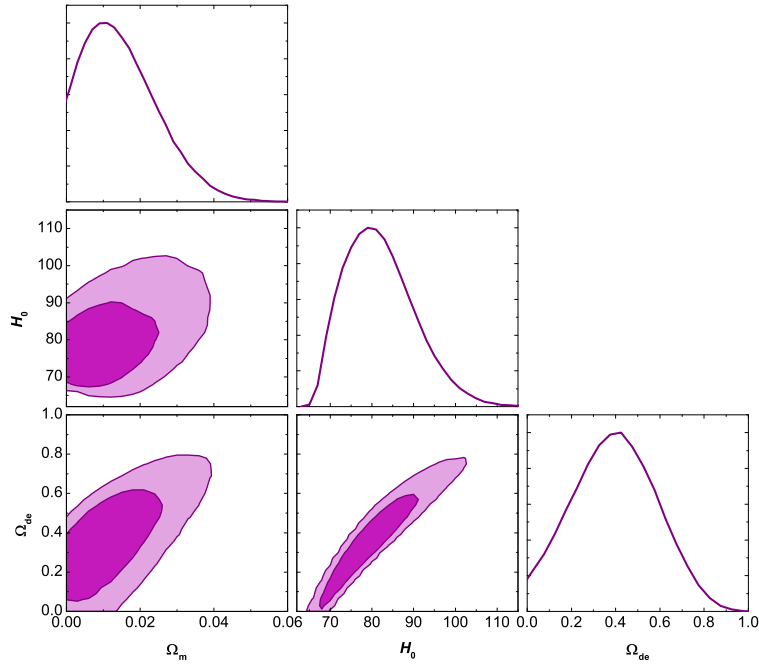


Fig. 8.— The 1-D probability distributions and 2-D regions with the 1σ and 2σ contours corresponding to the parameters Ω_m , Ω_{de} , and H_0 in the best-fit Λ CDM model, using the simulated sample with 350 ages, assuming $R_h = ct$ as the background cosmology.

sample with 350 ages in the $R_h = ct$ Universe are $\Omega_m = 0.011$, $\Omega_{de} = 0.37^{+0.25}_{-0.32}$ (1σ), and $H_0 = 79.9^{+10.7}_{-12.7}$ (1σ) $\text{km s}^{-1} \text{Mpc}^{-1}$. Note that the simulated ages provide a good constraint on Ω_{de} , but only a weak one on Ω_m ; only an upper limit of ~ 0.025 can be set at the 1σ confidence level.

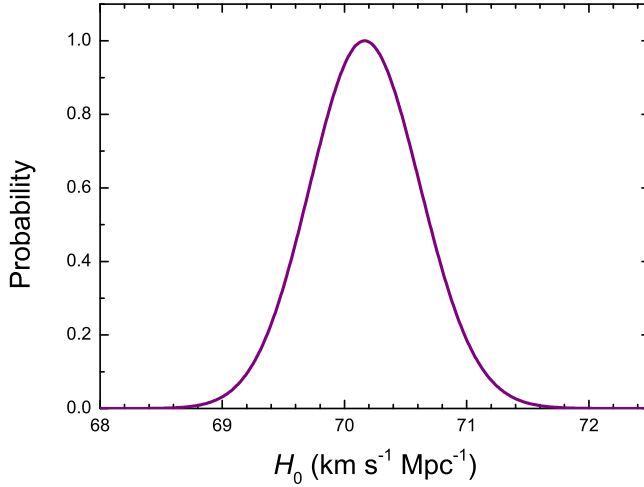


Fig. 9.— The 1-D probability distribution for the parameter H_0 in the $R_h = ct$ universe, using a sample of 350 ages, simulated with $R_h = ct$ as the background cosmology. The assumed value for H_0 in the simulation was $H_0 = 70 \text{ km s}^{-1} \text{Mpc}^{-1}$.

In Figure 9, we show the corresponding 1-D probability distribution of H_0 for the $R_h = ct$ universe. The best-fit value for the simulated sample is $H_0 = 70.2^{+0.5}_{-0.5}$ (1σ) $\text{km s}^{-1} \text{Mpc}^{-1}$. The assumed value for H_0 in the simulation was $H_0 = 70 \text{ km s}^{-1} \text{Mpc}^{-1}$.

Since the number N of data points in the sample is now much greater than one, the most appropriate information criterion to use is the BIC. The logarithmic penalty in this model selection tool strongly suppresses overfitting if N is large (the situation we have here, which is deep in the asymptotic regime). With $N = 350$, our analysis of the simulated sample shows that the BIC would favour the $R_h = ct$ Universe over ΛCDM by an overwhelming likelihood of 99.7% versus only 0.3% (i.e., the prescribed 3σ confidence limit).

6.2. Assuming Λ CDM as the Background Cosmology

In this case, we assume that the background cosmology is Λ CDM, and seek the minimum sample size to rule out $R_h = ct$ at the 3σ confidence level. We have found that a minimum of 45 galaxy ages are required to achieve this goal. To allow for the greatest flexibility in the Λ CDM fit, here too we relax the assumption of flatness, and allow Ω_{de} to be a free parameter, along with Ω_m . In Figure 10, we show the 1-D probability distribution for each parameter (Ω_m , Ω_{de} , H_0), and 2-D plots of the 1σ and 2σ confidence regions for two-parameter combinations. The best-fit values for Λ CDM using this simulated sample with 45 galaxy ages are $\Omega_m = 0.28^{+0.12}_{-0.11}$ (1σ), $\Omega_{de} = 0.30$, and $H_0 = 60.1^{+9.2}_{-7.6}$ (1σ) $\text{km s}^{-1} \text{Mpc}^{-1}$. Note that the simulated ages now give a good constraint on Ω_m , but only a weak one on Ω_{de} ; only an upper limit of ~ 0.83 can be set at the 1σ confidence level.

The corresponding 1-D probability distribution of H_0 for the $R_h = ct$ universe is shown in Figure 11. The best-fit value for the simulated sample is $H_0 = 85.0^{+1.6}_{-1.5}$ (1σ) $\text{km s}^{-1} \text{Mpc}^{-1}$. This is similar to that in the standard model, but not exactly the same, reaffirming the importance of reducing the data separately for each model being tested. With $N = 45$, our analysis of the simulated sample shows that in this case the BIC would favour Λ CDM over $R_h = ct$ by an overwhelming likelihood of 99.7% versus only 0.3% (i.e., the prescribed 3σ confidence limit).

These results were obtained assuming a uniform incubation time $\langle\tau\rangle$ throughout the mock sample. Of course, if the incubation time is distributed, the corresponding uncertainty in its distribution function will contribute to the variance of the “observed” age of the Universe. Thus, adding the scatter in the uncertain incubation time τ_c to the simulations would change the constructed sample size required to achieve the 3σ results discussed above. We have therefore carried out additional simulations using a Gaussian $P(\tau)$ distribution of the incubation time, and an assumed dispersion $\sigma_\tau = 0.3$. The corresponding uncertainty

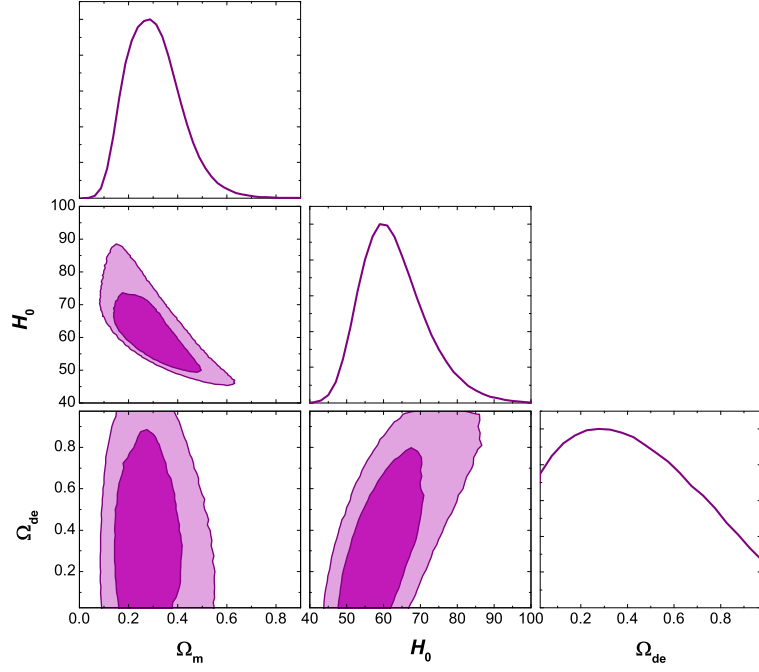


Fig. 10.— Same as Figure 8, except now with a flat Λ CDM as the (assumed) background cosmology. The simulated model parameters were $\Omega_m = 0.27$ and $H_0 = 70$ km s⁻¹ Mpc⁻¹.

in τ contributes to the variance of the “observed” age of the Universe. From these results, we estimate that a sample of about 55 galaxy ages would be needed to rule out $R_{\text{h}} = ct$ at a $\sim 99.7\%$ confidence level if the real cosmology were Λ CDM, while a sample of at least 500 ages would be needed to similarly rule out Λ CDM if the background cosmology were instead $R_{\text{h}} = ct$.

7. Discussion and Conclusions

In this paper, we have used the sample of high-redshift galaxies with measured ages to compare the predictions of several cosmological models. We have individually optimized the parameters in each case by minimizing the χ^2 statistic. Using a sample of 32 passively evolving galaxies distributed over the redshift interval $0.117 \leq z \leq 1.845$, we have demonstrated how these age-redshift data can constrain parameters, such as H_0 and Ω_{m} . For Λ CDM, these data are not good enough to improve upon the concordance values yet, but are approaching the probative levels seen with currently available gamma-ray burst luminosity data (Wei et al. 2013), strong gravitational-lensing measurements (see, e.g., Suyu et al. 2013), and measurements of the Hubble parameter as a function of redshift (Melia & Maier 2013).

Based solely on these 32 passively evolving galaxies, a comparison of the χ^2_{dof} for the $R_{\text{h}} = ct$ Universe and the concordance Λ CDM model shows that the age-redshift data do not yet favour either model. The $R_{\text{h}} = ct$ Universe fits the data with $\chi^2_{\text{dof}} = 0.435$ for a Hubble constant $H_0 = 67.2^{+4.5}_{-4.0} \text{ km s}^{-1} \text{ Mpc}^{-1}$ and a average delay time $\langle \tau \rangle = 2.72 \text{ Gyr}$. By comparison, the concordance model fits these same data with a reduced $\chi^2_{\text{dof}} = 0.523$, with a delay time $\langle \tau \rangle = 1.36 \text{ Gyr}$. Both are consistent with the view that none of these galaxies should have started forming prior to the transition from Population III to Population II stars at $\sim 300 \text{ Myr}$. However, if we relax some of the priors, and allow both Ω_{m} and H_0

to be optimized in Λ CDM, we obtain best-fit values $\Omega_m = 0.12_{-0.11}^{+0.54}$ and $H_0 = 94.3_{-35.8}^{+32.7}$ $\text{km s}^{-1} \text{Mpc}^{-1}$. The delay factor corresponding to this best fit is $\langle\tau\rangle = 1.62$ Gyr, with a $\chi^2_{\text{dof}} = 0.428$. The current sample favours $R_h = ct$ over the standard model with a likelihood of $\approx 66.5\% - 80.5\%$ versus $\approx 19.5\% - 33.5\%$.

We also analyzed the age-redshift relationship in cases where the delay factor τ may be different from galaxy to galaxy, and considered two representative distributions: (i) a Gaussian; and (ii) a top-hat. We found that the optimized cosmological parameters change quantitatively, though the qualitative results and conclusions remain the same, independent of what kind of the distribution one assumes for τ . Though one does not in reality expect the delay factor to be uniform, the fact that its distribution does not significantly affect the results can be useful for a qualitative assessment of the data. It also suggests that the outcome of our analysis is insensitive to the underlying assumptions we have made.

But though galaxy age estimates currently tend to slightly favour $R_h = ct$ over Λ CDM, the known sample of such measurements is still too small for us to completely rule out either model. We have therefore considered two synthetic samples with characteristics similar to those of the 32 known age measurements, one based on a Λ CDM background cosmology, the other on $R_h = ct$. From the analysis of these simulated ages, we have estimated that a sample of about 45 – 55 galaxy ages would be needed to rule out $R_h = ct$ at a $\sim 99.7\%$ confidence level if the real cosmology were in fact Λ CDM, while a sample of 350 – 500 ages would be needed to similarly rule out Λ CDM if the background cosmology were instead $R_h = ct$. These ranges allow for the possible contribution of an uncertainty in τ to the variance of the observed age of the Universe at each redshift. The difference in required sample size is due to Λ CDM’s greater flexibility in fitting the data, since it has a larger number of free parameters.

Both the Gaussian and Top-hat distributions that we have incorporated into this

study have assumed that the mean and scatter of the incubation time are constant with redshift. However, it would not be unreasonable to suppose that these quantities could have a systematic dependence on z . To examine how the results might change in this case, we have therefore also analyzed the real data using a Gaussian distribution

$$P(\tau) \propto \exp - \frac{[\tau - \tau_c \cdot (1 + z_i)^\alpha]^2}{2 [\sigma_\tau \cdot (1 + z_i)^\alpha]^2}, \quad (10)$$

with $\sigma_\tau = 0.3$ and, for simplicity, $\alpha = 1$. Figure 12(a) shows the corresponding distributions of τ_c and H_0 for the concordance Λ CDM model, with best-fit values $(H_0, \tau_c) = (82.5, 0.26)$. The analogous distributions for the $R_h = ct$ Universe are shown in Figure 12(b). In this case, the best fit corresponds to $(H_0, \tau_c) = (80.0, 0.73)$. The added redshift dependence has not changed the result that both models fit the passive galaxy age-redshift relationship comparably well, based solely on their reduced χ^2 's. Note that in this case, both the concordance Λ CDM and $R_h = ct$ models have the same free parameters (i.e., $(H_0$ and $\tau_c)$, so the information criteria should all provide the same results. Therefore, we show only the AIC results here. We find that the AIC does not favour either $R_h = ct$ or the concordance Λ CDM model, with relatively likelihoods of $\approx 54\%$ versus $\approx 46\%$. Note, however, that the best-fit value of τ_c in Λ CDM does not appear to be consistent with the supposition that all of the galaxies should have formed after the transition from Population III to Population II stars at $t \sim 300$ Myr.

An additional limitation of this type of work is the degree of uncertainty in the galaxy-age measurement itself. It is difficult to precisely constrain stellar ages for systems that are spatially resolved using stellar evolution models. We may be grossly underestimating how uncertain the age measurements of distant galaxies are. One ought to acknowledge this possibility and consider its impact on cosmological inferences. For example, redoing our analysis using a Gaussian $P(\tau)$ and an assumed dispersion $\sigma_\tau = 0.3$, but now with an additional 24% uncertainty on the age measurements, (i.e., twice as big

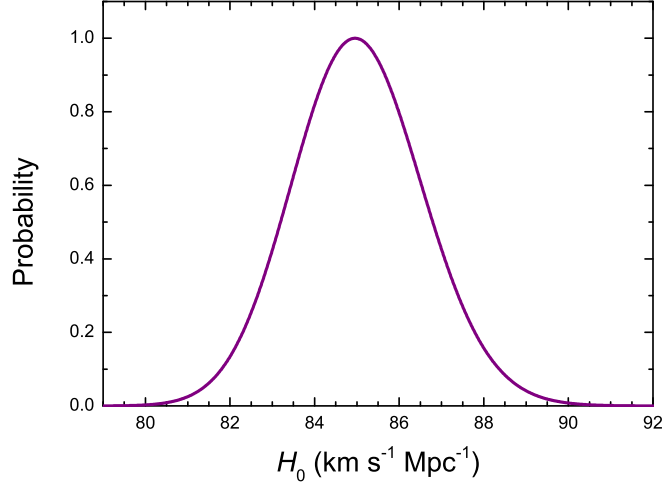


Fig. 11.— Same as Figure 9, except now with Λ CDM as the (assumed) background cosmology.

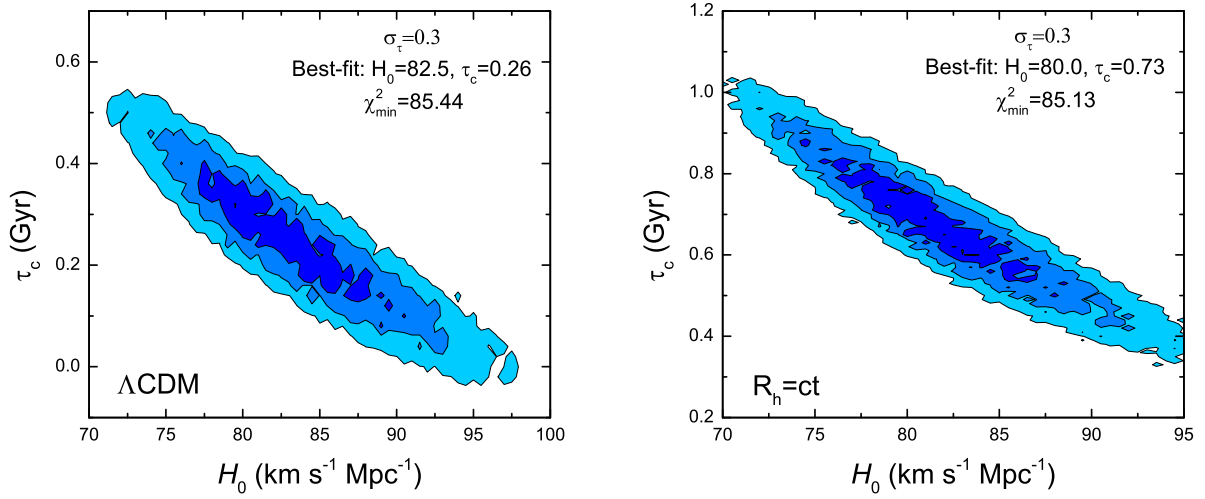


Fig. 12.— (a): Λ CDM with a Gaussian distribution for τ and a redshift dependent dispersion and mean value (see text). (b): Same as (a), except now for the $R_h = ct$ Universe.

as the value quoted in Dantas et al. 2009, 2011, and Samushia et al. 2010), produces the results shown in Figure 13. Panel (a) in this plot shows the corresponding distributions of τ_c and H_0 in the concordance Λ CDM model, with best-fit values $(H_0, \tau_c) = (77.5, 0.86)$. Not surprisingly, a comparison of Figure 13(a) with the left-middle panel in Figure 4 shows that, as the uncertainty in the age measurement increases, the constraints on model parameters weaken; nonetheless, the best-fit values of H_0 and τ_c are more or less the same.

The comparison between Λ CDM (Figure 13a) and $R_h = ct$ (Figure 13b) may be summarized as follows: the best-fit results are more or less the same for both the 12% and 24% uncertainties. Based solely on their minimum χ^2 values, both models fit the passive galaxy age-redshift relationship comparably well. The AIC does not favour either $R_h = ct$ or the concordance Λ CDM model, regardless of how uncertain the galaxy-age measurements are, with relative likelihoods of $\approx 51\%$ versus $\approx 49\%$.

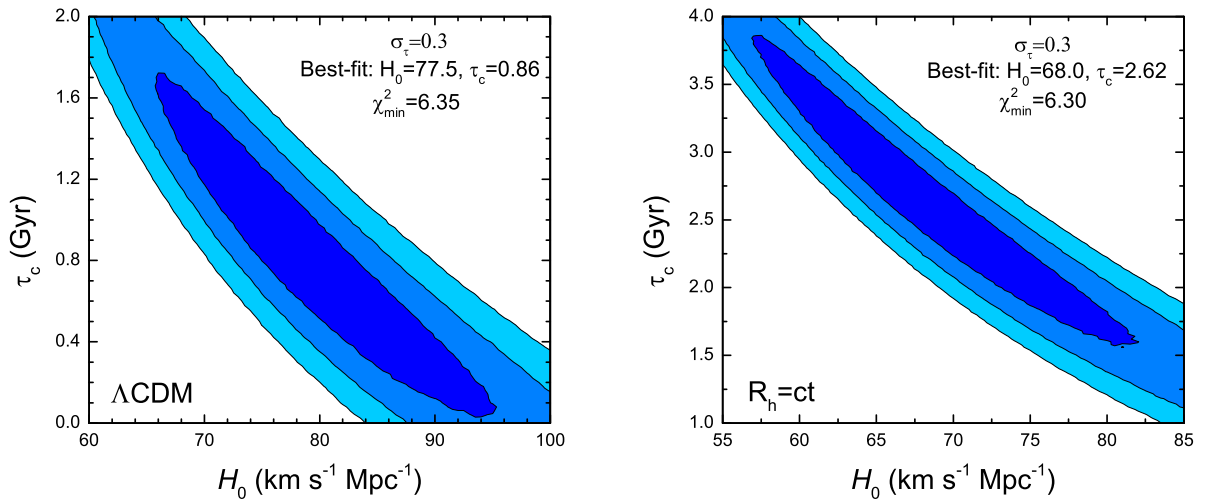


Fig. 13.— (a): Λ CDM with a Gaussian distribution of τ and an (assumed) uncertainty of 24% in the age measurement. (b): Same as (a), except now for the $R_h = ct$ Universe.

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